



Mn abundances in the stars of the Galactic disc with metallicities $-1.0 < [\text{Fe}/\text{H}] < 0.3$

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ABSTRACT

In this work, we present and discuss the observations of the Mn abundances for 247 FGK dwarfs, located in the Galactic disc with metallicity $-1 < [\text{Fe}/\text{H}] < +0.3$. The observed stars belong to the substructures of the Galaxy thick and thin disks, and to the Hercules stream. The observations were conducted using the 1.93 m telescope at Observatoire de Haute-Provence (OHP, France) equipped with the echelle-type spectrographs ELODIE and SOPHIE. The abundances were derived under the LTE approximation, with an average error for the $[\text{Mn}/\text{Fe}]$ ratio of 0.10 dex. For most of the stars in the sample, Mn abundances are not available in the literature. We obtain an evolution of $[\text{Mn}/\text{Fe}]$ ratio with the metallicity $[\text{Fe}/\text{H}]$ consistent with previous data compilations. In particular, within the metallicity range covered by our stellar sample, the $[\text{Mn}/\text{Fe}]$ ratio is increasing with the increase of metallicity. This is due to the contribution to the Galactic chemical evolution of Mn and Fe from thermonuclear supernovae. We confirm the baseline scenario where most of the Mn in the Galactic disc and in the Sun is made by thermonuclear supernovae. In particular, the effective contribution from core-collapse supernovae to the Mn in the Solar system is about 10–20 per cent. However, present uncertainties affecting the production of Mn and Fe in thermonuclear supernovae are limiting the constraining power of the observed $[\text{Mn}/\text{Fe}]$ trend in the Galactic discs on, e.g. the frequency of different thermonuclear supernovae populations. The different production of these two elements in different types of thermonuclear supernovae needs to be disentangled by the dependence of their relative production on the metallicity of the supernova progenitor.

Key words: stars: abundances – stars: late-type – galaxy: disc – galaxy: evolution.

1 INTRODUCTION

Manganese (Mn, $Z = 25$) is a monoisotopic element member of the iron group. In stellar spectra, several Mn absorption lines are known. Since early studies of stellar chemical composition, it was observed that in metal-poor stars, Mn has a different behaviour with respect to Fe compared to other iron-peak elements. (e.g. Wallerstein 1962). The chemical evolution of $[\text{Mn}/\text{Fe}]$ is also different compared to α -elements (O, Mg, Si, S, Ca and Ti), which abundances increase with the metallicity decreasing (e.g. Gratton 1989).

Today, a large number of Mn spectroscopic observations are available for stars with different age and metallicity and from different

galactic hosts, e.g. from our Galaxy, from globular clusters (GCs) including Omega-Cen, and from dwarf spheroidal galaxies (e.g. Prochaska et al. 2000; McWilliam, Rich & Smecker-Hane 2003; Alves-Brito, Barbuy & Allen 2007; Sobeck et al. 2006; Cunha et al. 2010; Pancino et al. 2011). In particular, the Mn abundance survey in GCs and field stars within the metallicity range $-2.7 < [\text{Fe}/\text{H}] < -0.7$ by Sobeck et al. (2006) found consistent average $[\text{Mn}/\text{Fe}]$ ratios ($[\text{Mn}/\text{Fe}] = -0.36$ and -0.37 , respectively) for those two populations. This is consistent with the fact that for metallicities $[\text{Fe}/\text{H}] \lesssim -1$ core-collapse supernovae (CCSNe) are the only astrophysical producers of Mn in the Galaxy (e.g. Thielemann, Nomoto & Hashimoto 1996; Woosley, Heger & Weaver 2002; Nomoto, Kobayashi & Tominaga 2013). The same is true for GCs, where the mass of the cluster is not large enough to keep the supernovae ejecta, and all Mn and Fe observed is due to pollution from massive

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stars before the GC formed. This is the main reason for the good agreement for the average $[\text{Mn}/\text{Fe}]$ observed in unevolved halo stars and in GCs (Sobeck et al. 2006). On the other hand, the star-to-star scatter reported by Sobeck et al. (2006) is $-0.6 \lesssim [\text{Mn}/\text{Fe}] \lesssim 0$, that is larger than the reported observational errors. The same conclusion may be derived by the observation in halo stars ($-1.1 \lesssim [\text{Mn}/\text{Fe}] \lesssim 0.5$) and from Omega Cen ($-0.8 < [\text{Mn}/\text{Fe}] < -0.2$; Cunha et al. 2010; Pancino et al. 2011). Such a large spread is difficult to reconcile with baseline one-dimensional CCSN models (e.g. Woosley & Weaver 1995; Limongi, Straniero & Chieffi 2000), where for an amount of ^{56}Ni ejected in the order of $0.1 M_{\odot}$ the stellar yields show an $[\text{Mn}/\text{Fe}]$ ratio $\gtrsim -0.5$. Therefore, the origin of the $[\text{Mn}/\text{Fe}]$ in the range $-1.0 \lesssim [\text{Mn}/\text{Fe}] \lesssim -0.5$ is not clearly understood (Andrievsky, Korotin & Martin 2007), although an increase of Y_e in the ejecta (see fig. 5 in Thielemann et al. 1996), as expected from neutrino interactions (see Fröhlich et al. 2006) can explain smaller values. An alternative scenario, in order to reproduce the low $[\text{Mn}/\text{Fe}]$ ratio is the contribution from hypernovae, ejecting large quantities of Fe compared to Mn (e.g. Umeda & Nomoto 2005; Nomoto et al. 2006, 2013).

The observation of Mn abundances may be affected by deviations from the LTE. Bergemann & Gehren (2008) have found that $[\text{Mn}/\text{Fe}]$ ratio measured in LTE approximation might underestimate the real Mn abundances up to 0.5–0.6 dex for the metal-poor stars, and ~ 0.1 dex for stars of solar-like metallicities.

For metallicities higher than $[\text{Fe}/\text{H}] \sim -1$ typical of the Galactic discs, the observation of the $[\text{Mn}/\text{Fe}]$ ratio is still controversial. Nissen et al. (2000) and Reddy, Lambert & Allende Prieto (2006) reported similar Mn abundance trends with $[\text{Fe}/\text{H}]$ for thick and thin disc stars. On the other hand, Feltzing, Fohlman & Bensby (2007) and Battistini & Bensby (2015) found a different behaviour in the two stellar populations. A similar result is obtained by Hawkins et al. (2015) for giant stars.

The nucleosynthesis of Mn becomes more complex for metallicities typical of the Galactic disc. Thermonuclear supernovae (SNIa; e.g. Hillebrandt et al. 2013, for a recent review) start to contribute to the production of Mn and Fe for $[\text{Fe}/\text{H}] \gtrsim -1$ (Matteucci & Greggio 1986), leading to the observed increasing trend of $[\text{Mn}/\text{Fe}]$ up to the present solar values. In particular, most of the Fe and Mn observed today in the Solar system are made by SNe Ia (e.g. Timmes, Woosley & Weaver 1995; Cescutti et al. 2008; Kobayashi, Karakas & Umeda 2011). The production of Fe in SNe Ia as ^{56}Ni is not completely independent from the initial metallicity of the star, and its production tends to decrease with the increasing metallicity (Timmes, Brown & Truran 2003; Travaglio, Hillebrandt & Reinecke 2005; Bravo et al. 2010). On the other hand, Mn production increases with the initial metallicity of the SNIa progenitor (e.g. Iwamoto et al. 1999; Travaglio et al. 2005), which is consistent with the increasing $[\text{Mn}/\text{Fe}]$ with $[\text{Fe}/\text{H}]$. The metallicity dependence of Mn production in SNe Ia was also inferred by Galactic chemical evolution (GCE) simulations (Cescutti et al. 2008).

One of the main uncertainties affecting the evolution of Mn abundance in the Galactic disc is our present understanding of SNIa nucleosynthesis. The two historical scenarios proposed for SNe Ia are the single-degenerate scenario (SDS), where an accreting white dwarf (WD) is reaching the Chandrasekhar mass by accreting material from a stellar binary companion, or the double-degenerate scenario (DDS), where an SNIa is formed by a merger of two CO WDs. Based on different motivations, in the last decades one or the other scenario have been favoured (see discussion in Hillebrandt et al. 2013). Another scenario related to the SDS is the double-detonation SN, where the SNIa explosion is triggered by the He

detonation initiated in the external He shell during the accretion on a sub-Chandrasekhar progenitor (Fink et al. 2010; Pakmor et al. 2010).

Based on GCE calculations of the $[\text{Mn}/\text{Fe}]$ ratio in the Galactic disc, Matteucci et al. (2009) proposed that two different types of SNIa contributors are needed to reproduce the observations. More recently, based on the use of stellar yields from multidimensional hydrodynamics simulations of SNe Ia, Seitenzahl et al. (2013b) confirmed that both SNIa scenarios with a mass close to the Chandrasekhar limit and with sub-Chandrasekhar mass (i.e. DDS or double-detonation in an accreting sub-Chandrasekhar progenitor) are needed to fit the $[\text{Mn}/\text{Fe}]$ evolution. In particular, the yields of the first ones carry a much larger Mn/Fe ratio than the second ones.

In this work, we provide the measurement for Mn abundances for 247 F–G–K-type dwarf stars located in the thin and thick disc populations, and for the Hercules stream, covering the metallicity range $-1.0 < [\text{Fe}/\text{H}] < 0.3$. The main purpose of this survey is to provide new constraints helping to definitely establish the origin of the Mn production in the Galactic disc. The paper is organized as follow. The observations and selection of stars, and definition of the main stellar parameters are described in Section 2. The abundance determinations for Mn and the error analysis are presented in Section 3. In Section 4 the results are compared with other measurements available in the literature. The implications of the results and the nucleosynthesis of Mn in stars is reported in Section 5. Conclusions are drawn in Section 6.

2 OBSERVATIONS, SELECTION AND PARAMETERS OF THE DISC STARS

The investigated stellar spectra were obtained using the 1.93 m telescope at Observatoire de Haute-Provence (OHP, France) equipped with echelle-type spectrographs, namely SOPHIE, resolving power $R = 75\,000$ (Perruchot et al. 2008) and ELODIE, $R = 42\,000$ (Baranne et al. 1996), in the wavelength range λ 4400–6800 Å and signal-to-noise ratio of about 100–300.

To distinguish between stars in the thin disc, thick disc and the Hercules stream, we considered the probability that a star belongs to either one of them accounting for its spatial velocity, kinematic disc parameters, and the quantity and percentage of every disc stars in our sample. The stars that belong to the Galactic substructures were selected using the technique described in Mishenina et al. (2004). The primary processing of spectra was carried out immediately during observations (Katz et al. 1998). Further spectra processing such as the continuum placement, line depth and equivalent width (EW) measurements, etc., was conducted using the DECH20 software package by Galazutdinov (1992). The atmospheric parameters of target stars has been determined earlier. The methods applied are described in detail in Mishenina & Kovtyukh (2001) and Mishenina et al. (2004, 2008).

The effective temperatures T_{eff} are defined by calibration of line-depth ratios ($R1/R2$) for spectral line pairs that markedly differ in their low-level excitation potential Kovtyukh et al. (2003). A large number of calibrations (80–103) permitted to reduce the influence of errors in the line-depth and atmospheric parameter measurements on the resulting temperature estimates. The intrinsic accuracy of the method applied for dwarfs is 5–45 K. For the stars with $[\text{Fe}/\text{H}] < -0.5$ the T_{eff} were estimated by adjustment of far-wings of the H α line (Mishenina & Kovtyukh 2001).

The surface gravities $\log g$ values are computed by the ionization balance of the neutral and ionized iron. This method implies that similar abundances are obtained from the neutral iron Fe I and

Table 1. Atmospheric parameters and Mn abundance $\log A(\text{Mn})$ each used line (λ , A) and $[\text{Mn}/\text{Fe}]$ ratio for our target stars.

Star	$T_{\text{eff},\text{K}}$	$\log g$	$[\text{Fe}/\text{H}]$	V_t	4783 Å	4823 Å	5432 Å	6013 Å	6021 Å	$[\text{Mn}/\text{Fe}]$
Sun					5.30	5.25	5.27	5.24	5.24	
Thick disc										
HD 245	5400	3.4	-0.84	0.7	4.32	4.35	4.10	4.18	4.18	-0.16
HD 3765	5079	4.3	0.01	1.1			5.30	5.30	5.30	0.04
HD 6582	5240	4.3	-0.94	0.7	4.20	4.20	4.18	4.15	4.15	-0.14
HD 13783	5350	4.1	-0.75	1.1	4.35	4.40	4.40		4.25	-0.17
HD 18757	5741	4.3	-0.25	1.0	4.95	4.95	4.95	4.85	4.85	-0.10
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...
...

ionized iron Fe II lines. Its accuracy is affected by a number of factors, such as uncertainty of oscillator strengths of the $\log gf$ lines and thermal structure of atmospheric models, and NLTE effects.

However, the determination of the surface gravity by the parallax method is also affected by uncertainties. In order to determine the stellar mass by using theoretical evolution tracks, it is necessary to measure in advance the metallicity and α -element enrichment, leading to an uncertainty in the mass determination of $\sim 0.2 M_{\odot}$, which corresponds to an error for the surface gravity of ~ 0.1 dex. Therefore, NLTE effects and atmospheric model uncertainties still affect the analysis. As reported by Allende Prieto et al. (1999), the astrometric and spectroscopic methods provide consistent results in the metallicity range $-1 < [\text{Fe}/\text{H}] < 0.3$. Having compared the resulting $\log g$ obtained by us using the ionization balance method with the $\log g$ value computed by using the parallax in Allende Prieto et al. (1999) for 39 stars in common, leads to differences not exceeding 0.1 dex on average (Mishenina et al. 2004).

The microturbulence velocity V_t is derived considering that the iron abundance obtained from the given Fe I line is not correlated with the EW of that line. The adopted value of metallicity $[\text{Fe}/\text{H}]$ is the iron abundance derived from the Fe I lines. The determination errors are: for the effective temperatures $\delta T_{\text{eff}} = \pm 100$ K, the surface gravities $\delta \log g = \pm 0.2$ dex, the microturbulence velocity $\delta V_t = \pm 0.2$ km s $^{-1}$. The obtained parameter values and their comparison with the results of other authors are reported in Mishenina et al. (2004, 2013).

3 THE Mn ABUNDANCE

The Mn abundances were derived by computing the synthetic spectrum in the region of the Mn lines by the newly updated STARS LTE software (Tsymbal 1996) using the Kurucz models (Castelli & Kurucz 2004). The Mn I lines undergo hyperfine-structure (HFS) splitting, due to the interaction of the magnetic moment of the nucleus with the magnetic field of the electrons. The list of lines and HFS data were taken from Prochaska et al. (2000). The van der Waals damping constant C_6 was taken from Bergemann & Gehren (2008). Atomic data for other lines required to compute the synthetic spectrum in the region of the Mn I lines were taken from the Vienna Atomic Line Database VALD (Kupka et al. 1999). The solar Mn I abundances are derived for each line by the solar spectra reflected by the moon and asteroids. Two echelle-type spectrographs were adopted. In order to select reliable fitting lines, we started with derivation of the Mn abundances from 16 Mn I lines in the atmospheres of the Sun and some other stars, using spectra obtained with both spectrographs. A total of six solar spectra were used, including three spectra obtained with the echelle-type spectrograph ELODIE

and three spectra received by the echelle-type spectrograph SOPHIE. Five lines were obtained: 4783, 4823, 5432, 6013 and 6021 Å. The EWs averaged by spectra for each spectrograph differ by 1–2 per cent for strong lines and 2–3 per cent for the 5432 Å line. The mean difference of EWs of Mn lines in spectra obtained with two spectrographs is $\langle (\text{EWs}(\text{Mn})_{\text{SOPHIE}} - \text{EWs}(\text{Mn})_{\text{ELODIE}}) \rangle = -1.8 \pm 0.7$ mÅ. Further, we also compared of EWs of Fe I lines measured in solar spectra obtained with two spectrographs. The mean difference in this case is $\langle (\text{EWs}(\text{Fe})_{\text{SOPHIE}} - \text{EWs}(\text{Fe})_{\text{ELODIE}}) \rangle = -1.2 \pm 2.3$ mÅ. Hence, the Mn I abundance derived from selected lines in the solar spectra obtained with both spectrographs are consistent. Lines 4783 and 4823 Å are rather strong, so they were not used in the analysis of cooler stars ($T_{\text{eff}} < 5200$ K) and those more enriched in iron ($[\text{Fe}/\text{H}] > 0$). For each line, we derived the Mn abundances in the solar and stellar spectra, by fitting locally the observed spectra with the synthetic model. For a given line, we compare the abundance value to the solar abundance. We adopted this differential approach to eliminate the impact of potential errors in the oscillator strengths led to the following adopted atomic parameters. Finally, we derived the stellar Mn abundance by averaging the values obtained for each line (Table 1). The mean solar abundance computed for the lines with $\log gf$ from VALD data base (Kupka et al. 1999) and the solar model from Castelli & Kurucz (2004) and adopted values for each line in this work are 5.3, 5.25, 5.27, 5.24 and 5.24 for the lines λ 4783, 4823, 5432, 6013, 6021 Å, respectively. In Fig. 1 we show a few examples of profile fitting for Mn lines.

For a sample of stars also analysed by other studies, we compare the inferred atmospheric parameters in Table 2. Impact of these variations are within the observational errors for Mn abundance. For a detailed comparison, we show in Table 3 the stellar data for each star together with other works. Takeda (2007) and Nissen & Schuster (2011) show the largest variations compared to our T_{eff} , but they also report larger errors ($\sigma \gtrsim 100$ K). The biggest departure is for HD 4307 by Takeda (2007), with a difference for T_{eff} larger than 200 K. In the Table 4, we presented the comparison of the data of Battistini & Bensby (2015) and our determinations of parameters and Mn abundance for common stars. A good agreement is obtained between these two data sets.

3.1 Errors in abundance determinations

To determine the systematic errors in the Mn abundance resulting from uncertainties in the atmospheric parameter determinations, we derived the Mn abundance for several models with modified parameters ($\delta T_{\text{eff}} = \pm 100$ K, $\delta \log g = \pm 0.2$, $\delta V_t = \pm 0.1$). The Mn abundance variations with the modified parameters and the fitting

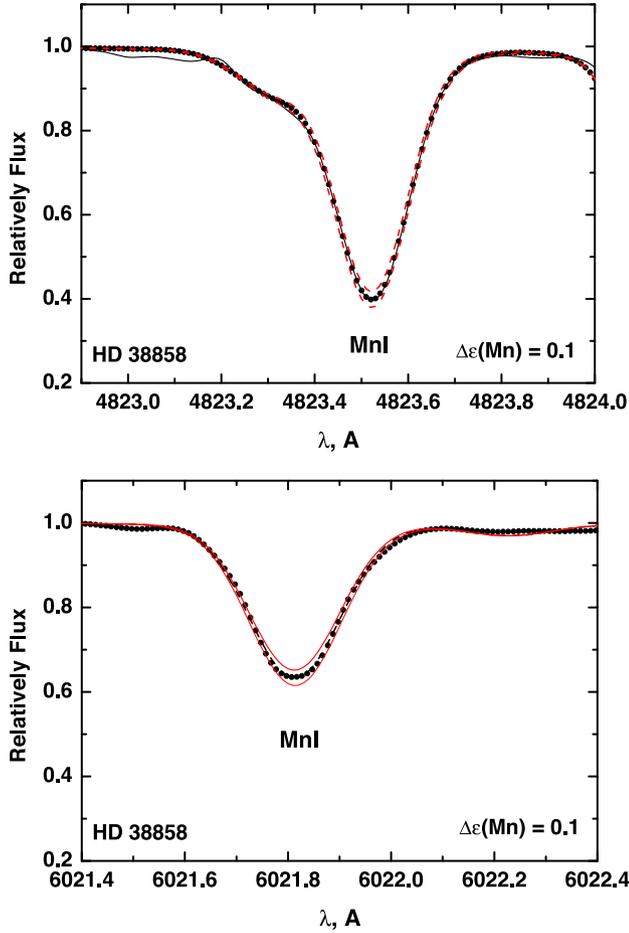


Figure 1. Example of the fitting of the observed spectrum (black dots with solid line) by the synthetic spectrum (red dotted line) in the area of the Mn lines.

errors for the computed and observed spectral line profiles (0.03 dex), are given in Table 5. The largest error occurs when the T_{eff} are determined inaccurately. Errors caused by uncertainties from other parameters are marginal. The total error associated with the determination of the Mn abundances is 0.08–0.10 dex (Table 5).

The graph of plotted points corresponding to the Mn abundances for each line in spectra of all investigated stars is presented in Fig. 2. In particular, there is no systematic difference in abundances obtained for a given line. The correlation between $\sigma([\text{Mn}/\text{H}])$ and $[\text{Mn}/\text{H}]$, where σ is the standard deviation, the mean value of which is 0.03, is shown in Fig. 2. No trend is observed on the graph of $\sigma[\text{Mn}/\text{H}]$ dependence on T_{eff} , $\log g$ or $[\text{Fe}/\text{H}]$ either (Fig. 2).

Table 2. Comparison of our parameters and Mn abundance determinations with the results of other authors for the n stars shared with our stellar sample.

	$\Delta(T_{\text{eff}})$	$\Delta(\log g)$	$\Delta([\text{Fe}/\text{H}])$	$\Delta([\text{Mn}/\text{Fe}])$	n
Feltzing et al. (2007)	24 ± 76	-0.03 ± 0.13	-0.01 ± 0.08	0.02 ± 0.06	10
Reddy et al. (2006)	105 ± 100	-0.18 ± 0.21	0.01 ± 0.11	0.14 ± 0.06	9
Takeda et al. (2007)	-14 ± 119	-0.06 ± 0.21	-0.04 ± 0.10	0.02 ± 0.08	31
Nissen et al. (2011)	7 ± 143	-0.03 ± 0.20	-0.05 ± 0.10	0.02 ± 0.02	4
Adibekyan et al. (2012)	28 ± 57	-0.07 ± 0.14	0.01 ± 0.04	0.01 ± 0.05	9
Battistini & Bensby (2015)	-4 ± 106	-0.10 ± 0.15	-0.03 ± 0.06	0.02 ± 0.06	22

Table 3. Comparison of atmospheric parameters and Mn abundance for common stars.

HD	Sources	T_{eff}	$\log g$	$[\text{Fe}/\text{H}]$	$[\text{Mn}/\text{Fe}]$
4307	Adibekyan et al. (2012)	5812	4.10	-0.23	-0.07
	Takeda (2007)	5648	3.75	-0.29	-0.09
	Our	5889	4.00	-0.18	-0.06
6582	Reddy et al. (2006)	5300	4.67	-0.86	-0.23
	Takeda (2007)	5330	4.54	-0.81	-0.12
	Our	5240	4.30	-0.94	-0.14
22879	Adibekyan et al. (2012)	5884	4.52	-0.82	-0.30
	Feltzing et al. (2007)	5920	4.33	-0.84	-0.18
	Nissen & Shuster (2011)	5759	4.25	-0.85	-0.27
	Reddy et al. (2006)	5722	4.50	-0.87	-0.39
	Our	5972	4.50	-0.77	-0.22
76932	Feltzing et al. (2007)	5875	4.10	-0.91	-0.23
	Nissen & Shuster (2011)	5877	4.13	-0.87	-0.25
	Reddy et al. (2006)	5783	4.09	-0.86	-0.35
	Our	5840	4.00	-0.95	-0.25
106516	Nissen & Shuster (2011)	6196	4.42	-0.68	-0.23
	Reddy et al. (2006)	6069	4.44	-0.74	-0.35
	Our	6165	4.40	-0.72	-0.23
125184	Adibekyan et al. (2012)	5680	4.10	0.27	0.06
	Takeda (2007)	5629	4.02	0.25	0.24
	Our	5695	4.30	0.31	0.02
157214	Reddy et al. (2006)	5605	4.49	-0.41	-0.24
	Takeda (2007)	5693	4.21	-0.37	-0.15
	Our	5820	4.50	-0.29	-0.11
159482	Nissen & Shuster (2011)	5737	4.31	-0.73	-0.23
	Reddy et al. (2006)	5630	4.58	-0.70	-0.32
	Our	5620	4.10	-0.89	-0.21
199960	Adibekyan et al. (2012)	5973	4.39	0.28	0.10
	Feltzing et al. (2007)	5924	4.26	0.28	0.23
	Takeda (2007)	5924	4.26	0.28	-0.04
	Our	5878	4.20	0.23	0.02
217014	Feltzing et al. (2007)	5789	4.34	0.20	0.02
	Takeda (2007)	5779	4.31	0.20	-0.09
	Our	5763	4.30	0.17	0.00

3.2 Analysis of the Mn spectral line parameters and evaluation of the effects of deviations from LTE on determination of the Mn abundance

In order to analyse and compare consistently the Mn abundance obtained in different works, it is necessary to take into account that the authors used different oscillator strengths $\log gf$ and different atomic data to account for HFS in their analysis. The values of the NLTE corrections for the Mn lines, including those used in our study, are investigated by Bergemann & Gehren (2007, 2008). The difference in the NLTE corrections change for different lines and depends on temperature and metallicity (Bergemann & Gehren 2008). The obtained variations for lines of various multiplets often exceed 0.10 dex. Therefore, the LTE Mn abundances obtained from

Table 4. Comparison of atmospheric parameters and Mn abundance for common stars with Battistini & Bensby (2015; hereafter BB 2015).

HIP	HD	T_{eff} (BB 2015)	$\log g$	[Fe/H]	[Mn/Fe]	T_{eff} our	$\log g$	[Fe/H]	[Mn/Fe]
6653	8648	5841	4.3	0.22	0.02	5790	4.2	0.12	0.02
16852	22484	6036	4.1	-0.03	-0.07	6037	4.1	-0.03	-0.04
17147	22879	5970	4.5	-0.81	-0.3	5972	4.5	-0.77	-0.22
22263	30495	5790	4.5	0.02	-0.18	5820	4.4	-0.05	-0.03
30545	45067	6042	3.9	-0.03	-	6058	4	-0.02	-0.09
38625	64606	5188	4.4	-0.91	-0.08	5250	4.2	-0.91	-0.11
38750	64815	5763	3.9	-0.35	-0.19	5864	4	-0.33	-0.14
44075	76932	5937	4.2	-0.9	-	5840	4	-0.95	-0.24
64792	115383	6185	4.3	0.25	0.02	6012	4.3	0.11	0.00
74537	135204	5200	4.4	-0.19	-	5413	4	-0.16	-0.03
81300	149661	5216	4.6	-0.01	-	5294	4.5	-0.04	-0.03
82588	152391	5322	4.5	-0.08	-	5495	4.3	-0.08	0.05
84905	157089	5915	4.3	-0.5	-	5785	4	-0.56	-0.19
86013	159482	5760	4.3	-0.81	-0.18	5620	4.1	-0.89	-0.21
86193	159909	5671	4.3	0.03	-0.07	5749	4.1	0.06	0.00
88622	165401	5794	4.5	-0.4	-0.16	5877	4.3	-0.36	-0.09
93966	178428	5656	4.2	0.15	-0.04	5695	4.4	0.14	0.00
97779	187897	5944	4.5	0.12	-0.03	5887	4.3	0.08	-0.02
98767	190360	5572	4.5	0.26	-	5606	4.4	0.12	0.05
103682	199960	6023	4.4	0.33	0.08	5878	4.2	0.23	0.02
104659	201891	5973	4.3	-1.08	-0.26	5850	4.4	-0.96	-0.28
113357	217014	5858	4.4	0.24	0.02	5763	4.3	0.17	0.00

Table 5. Abundance uncertainties due to atmospheric parameters.

Mn I lines	$\Delta T_{\text{eff}+}$	$\Delta \log g+$	$\Delta V_{\text{t}+}$	Tot+
	HD 22879	(5972/4.5/1.1/-0.77)		
4783	-0.08	0.02	0.03	
4823	-0.08	0.01	0.02	
5432	-	-	-	
6013	-0.06	0.02	0.02	
6021	-0.06	0.02	0.01	
Average	-0.07	0.017	0.02	0.08
	HD 26923	(5920/4.4/1.0/-0.03)		
4783	-0.06	0.03	0.03	
4823	-0.07	0.01	0.02	
5432	-0.08	0.02	0.03	
6013	-0.05	0.02	0.03	
6021	-0.06	0.02	0.01	
Average	-0.06	0.02	0.02	0.08
	HD 4635	(5103/4.4/0.8/0.07)		
4783	-0.10	0.04	0.02	
4823	-0.10	0.05	0.02	
5432	-0.08	0.03	0.06	
6013	-0.06	0.05	0.02	
6021	-0.04	0.05	0.04	
Average	-0.08	0.04	0.03	0.10

the lines of different multiplets must show systematic variations between them. According to Bergemann & Gehren (2008) due to the NLTE effects we should have obtained the systematic difference in LTE Mn abundance for two line groups of different multiplets. We used pairs of the lines 4783–4823 Å (multiplet 16, $E_{\text{low}} = 2.3$ eV) and 6013–6021 Å (multiplet 27, $E_{\text{low}} = 3.07$ eV). In Figs 3 and 4 we show the dependences of the difference in abundances $\Delta \log A = \log \text{Mn}_{4783} - \log \text{Mn}_{6013}$ on the effective temperature and metallicity for the thin and thick disc stars. There is no systematic trend observed. The average variations are 0.01 ± 0.04 and 0.02 ± 0.04 for the thin and thick discs, respectively. Thus, the observations do not

support the values of the LTE deviations obtained by Bergemann & Gehren (2008) at the given temperatures and metallicities.

The development of an adequate model of Mn atoms to account for the effects of deviations from LTE is complicated by the absence of detailed computations of atomic data, such as photoionization cross-section or parameters of radiative and shock transitions. The use of approximations such as an H-like approximation for Mn atoms, yields NLTE corrections that are not robust. Taking all this into account, we believe instead that the LTE determinations for the Mn abundance are correct within the given uncertainty of 0.1 dex.

4 RESULTS AND COMPARISON WITH THE LITERATURE

The Mn abundances obtained for our stellar sample is shown in Fig. 5, upper panel. In Fig. 6 our results are compared with other works for stars in the thin disc (Nissen et al. 2000; Gilli et al. 2006; Reddy et al. 2006; Feltzing et al. 2007; Takeda 2007; Adibekyan et al. 2012), in the thick disc (Reddy et al. 2006; Feltzing et al. 2007; Nissen & Schuster 2011; Adibekyan et al. 2012; Ishigaki, Aoki & Chiba 2013), in the Galactic halo (Preston & Sneden 2000; Cayrel et al. 2004; Hollek et al. 2011; Cohen et al. 2013; Ishigaki et al. 2013; Yong et al. 2013) and for different populations (Sobeck et al. 2006). Part of the observational scatter is due to the use of different Mn lines and different methods for the analysis in the papers presented above. Different works adopted the LTE approximation, while Bergemann & Gehren (2008) and BB 2015 used the NLTE approach. As we have shown above, our measurements for [Mn/Fe] are only marginally affected by the LTE assumption. Therefore, we are confident that we can also compare our results with these works. Most of the authors observe a similar increasing [Mn/Fe] trends with the increasing of the metallicity in the Galactic disc for $-1 \lesssim [\text{Fe}/\text{H}] \lesssim 0.3$. Takeda (2007) did not observe any clear trend, but they used only one Mn line (5040 Å). Stars from Gilli et al. (2006) show a larger scatter at near-to-solar metallicity compared to other works: in particular,

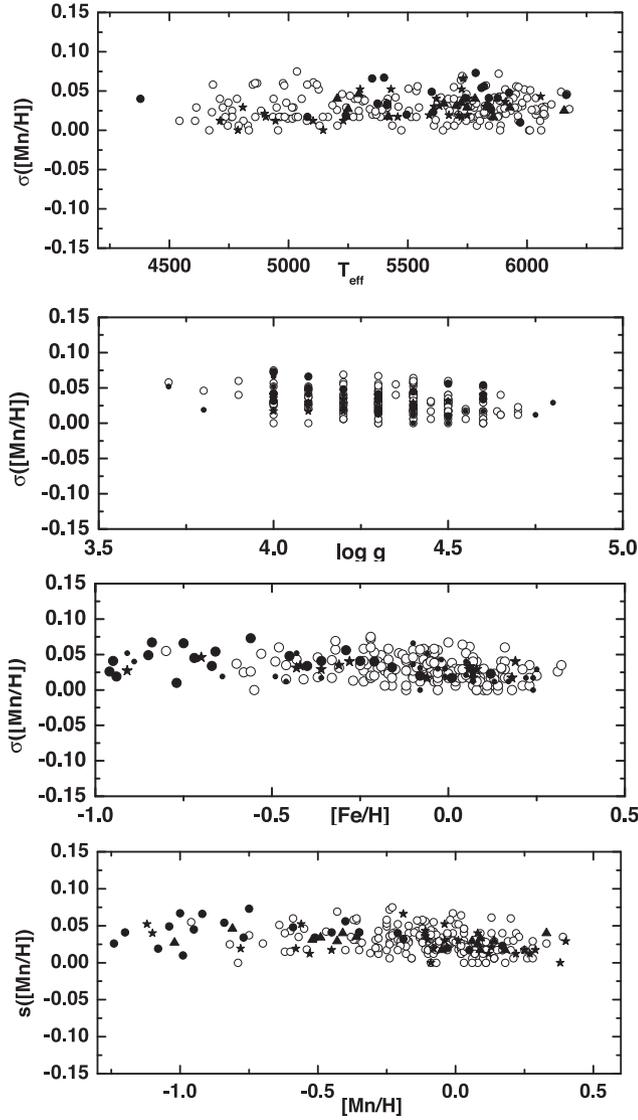


Figure 2. Dependences of σ [Mn/H] on atmospheric parameters, [Fe/H] and [Mn/H].

they obtain that stars hosting planetary systems show on average a larger Mn enrichment compared to stars without known planets.

The Mn abundances obtained for the thick and thin disc stellar populations in our stellar sample are comparable with the results by Reddy et al. (2006) with no difference between the thin disc and the thick disc within uncertainties. Stars in our sample belonging to the Hercules stream show abundances consistent with the other galactic disc populations. A detailed comparison between our results and the literature is given in Tables 3 and 4. In particular, for stars in common with Feltzing et al. (2007) we obtain similar results. The only exception is HD 199960, where we obtain a [Mn/Fe] lower by 0.21 dex that is within the errors of the two works. Note that the difference in the [Mn/Fe] for this star of our determination and the work of BB 2015 is 0.06 dex. For other stars with Mn included in our sample and also measured by BB 2015, the difference in [Mn/Fe] is within the determination error (0.10 dex) but for one star, HD 30495. In this case the difference is 0.15 dex, that is still consistent with our results within the errors given in BB 2015.

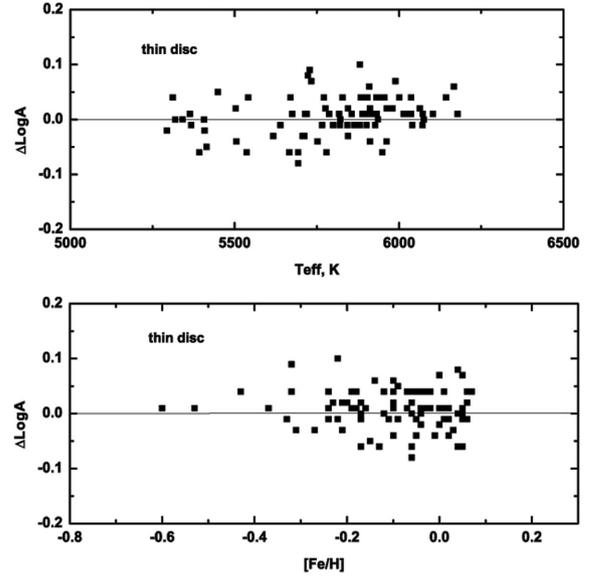


Figure 3. The dependence of the difference between Mn abundances determined with the Mn I lines at 4783 and 6013 Å for thin disc stars.

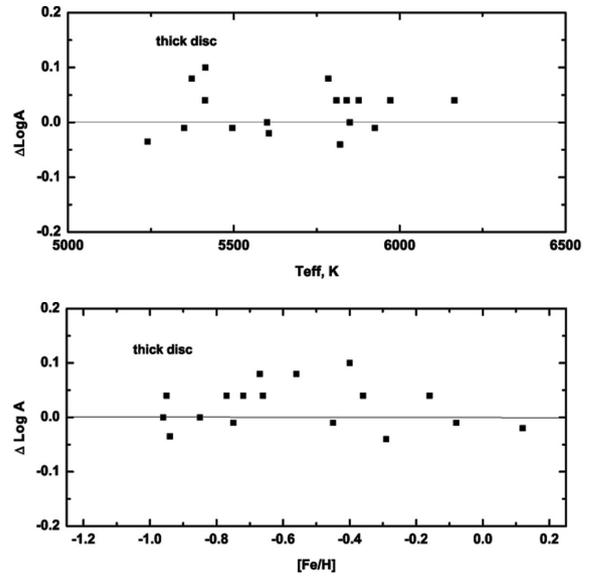


Figure 4. The dependence of the difference between Mn abundances determined with the Mn I lines at 4783 and 6013 Å for thick disc stars.

5 Mn OBSERVATIONS AND DISCUSSION: IMPLICATIONS FROM THE NUCLEOSYNTHESIS OF Mn IN STARS

In Fig. 6 we have shown the evolution of Mn compared to Fe for stars at different metallicities, ranging from low-metallicity stars to supersolar metallicities. For stars with $[\text{Fe}/\text{H}] \lesssim -1$, the abundance signature of Mn and Fe and therefore the observed [Mn/Fe] ratio is dominated by CCSNe (e.g. Kobayashi et al. 2011), while for higher metallicities the influence of SNe Type Ia dominates. The only stable isotope of Mn (^{55}Mn) is made in explosive (complete and incomplete) Si-burning conditions as unstable ^{55}Co , which decays afterwards via ^{55}Fe to ^{55}Mn . In complete Si-burning, this production occurs under normal freeze-out conditions only for sufficiently high densities and/or low entropies (see e.g. Thielemann, Nomoto & Yokoi 1986; Thielemann, Hashimoto & Nomoto 1990).

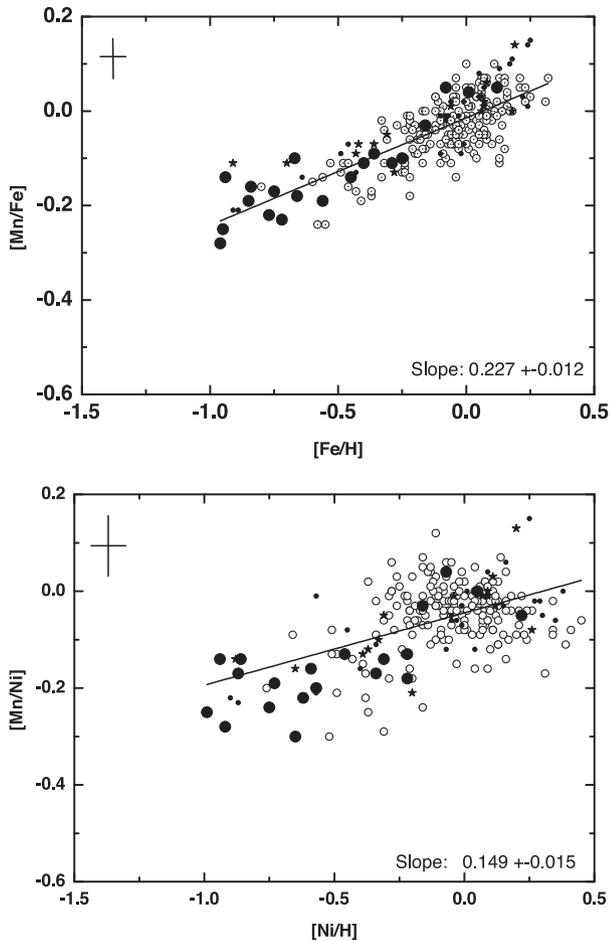


Figure 5. The trend of [Mn/Fe] ratio with respect to [Fe/H] is shown for our stellar sample: thin disc stars are marked as open symbols, thick disc stars as full symbols. Our determination for Hercules stream stars and unclassified stars are marked as black asterisks and points, respectively. In the bottom figure, we show the trend of [Mn/Ni] ratio with respect to [Ni/H].

Such conditions exist only in (near-)Chandrasekhar mass SNe Type Ia (Thielemann et al. 1986; Iwamoto et al. 1999; Brachwitz et al. 2000; Seitenzahl et al. 2009, 2013a; Fink et al. 2014). In the α -rich regime of complete Si-burning (the only type of complete Si-burning experienced in CCSNe) ^{55}Co abundance is moved over to ^{59}Cu which decays via ^{59}Ni to ^{59}Co . However, ^{55}Co is also produced in incomplete Si-burning, which takes place in CCSNe as well as SNe Ia (Thielemann et al. 1986, 1996; Woosley & Weaver 1995; Iwamoto et al. 1999; Nakamura et al. 1999; Brachwitz et al. 2000; Woosley et al. 2002). In all of these conditions, the production of ^{55}Co depends also on the electron fraction (Y_e) of the matter experiencing explosive burning. Y_e is the number of electrons per all nucleons (free and bound in nuclei), or the ratio of the number of all protons over all nucleons (i.e. neutrons plus protons). Thus, $Y_e = 0.5$ indicates a stellar composition with equal numbers of neutrons and protons, $Y_e < 0.5$ means that it is neutron rich and $Y_e > 0.5$ that is proton rich. In stellar evolution, the electron fraction Y_e first changes during H- and He-burning. Only marginal variations occur during the following C-burning and Ne-burning evolutionary stages. In advanced O-burning and then in Si-burning stages, the Y_e decreases significantly (e.g. Thielemann & Arnett 1985). However, according to present theoretical stellar models these last regions will not be

ejected by the SN explosion, or will not host the thermodynamic conditions needed to make Mn during the explosion.

There exists, however, a Y_e change in stellar models as a function of metallicity. In H-burning, the CNO-isotopes are burned essentially to ^{14}N (an $N = Z$ nucleus) which is moved in He-burning to ^{22}Ne (an $N = Z+2$ nucleus due to the beta-decay of ^{18}F to ^{18}O during the α -capture chain based on ^{14}N). In this way the metallicity (given predominantly by CNO abundances) is turned into the abundance of ^{22}Ne , which differs from the $Z/A = 0.5$ of the other He-burning products and determines the electron fractions $Y_e = \sum_i Z_i Y_i$. For lowest metallicities ($[\text{Fe}/\text{H}] = -\infty$) this relates to $Y_e = 0.5$ after He-burning, for solar metallicities ($[\text{Fe}/\text{H}] = 0$) to $Y_e = 0.499$, for supersolar metallicities ($[\text{Fe}/\text{H}] = 0.25, 0.5$) to $Y_e = 0.498, 0.496$. ^{55}Co , the radioactive progenitor of ^{55}Mn , has a $Z/A = 27/55$ of 0.491, which would be the Y_e value at which the highest ^{55}Mn production is expected, while low-metallicity stars would not produce ^{55}Mn if the Y_e is only determined by initial metallicity.

5.1 Massive stars

In massive stars, as we mentioned above, Mn is made mostly by the SN explosion in incomplete explosive Si-burning conditions as ^{55}Co (e.g. Woosley & Weaver 1995; Thielemann et al. 1996; Nakamura et al. 1999; Woosley et al. 2002). Mn production is increasing with the increase of the initial metallicity or, in other words, with the decrease of the electron fractions Y_e (e.g. Thielemann et al. 1996; Nakamura et al. 1999; Nomoto et al. 2013). Most of Fe is made as radiogenic ^{56}Fe from radioactive ^{56}Ni . This isotope is mainly made in complete explosive Si-burning conditions as primary product (i.e. independent of the initial metallicity of the star). However, a significant fraction is produced together with ^{55}Co in less-extreme (incomplete Si-burning) SN conditions, but with an increasing production with increasing Y_e (or decreasing of the initial metallicity), which is the opposite compared to ^{55}Co . This scenario is becoming more complicated once the zoo of different CCSN types is considered. For instance, hypernovae tend to produce lower [Mn/Fe] ratios compared to the less-energetic SN Type II, due to a larger production of ^{56}Ni and more extended α -rich freeze-out zones of complete Si-burning in comparison to incomplete Si-burning (e.g. Nakamura et al. 1999, 2001; Maeda et al. 2002). Asymmetries before and after the CCSN, and the SN-shock propagation through the massive star progenitor (e.g. Thielemann et al. 2011; Hix et al. 2014; Wongwathanarat, Mueller & Janka 2015) will affect the final Mn/Fe ratio in the SN ejecta, possibly explaining the large [Mn/Fe] spread observed in the early Galaxy (Fig. 6). Within this observational scatter, most of halo stars show an $[\text{Mn}/\text{Fe}] \sim -0.4$. Considering that CCSNe are producing about 30–50 per cent of the Fe observed in the Solar system, this means that 12–20 per cent of the solar Mn is made by CCSNe.

5.2 Type Ia supernovae

For $[\text{Fe}/\text{H}] \gtrsim -1$, the [Mn/Fe] ratio in disc stars is increasing up to the solar ratio, due to the contribution from SNe Ia (SNIa; e.g. Hillebrandt et al. 2013) with on average $[\text{Mn}/\text{Fe}] > 0$. This is due to the fact that (near-)Chandrasekhar mass SNe Type Ia experience also normal freeze-out conditions from nuclear statistical equilibrium in complete, explosive Si-burning, caused by high central densities and low entropies (see e.g. Thielemann et al. 1986; Iwamoto et al. 1999; Brachwitz et al. 2000; Seitenzahl et al. 2009, 2013a; Fink et al. 2014). In these inner high-density regions, the Y_e is not due to the initial stellar metallicity (see introduction to

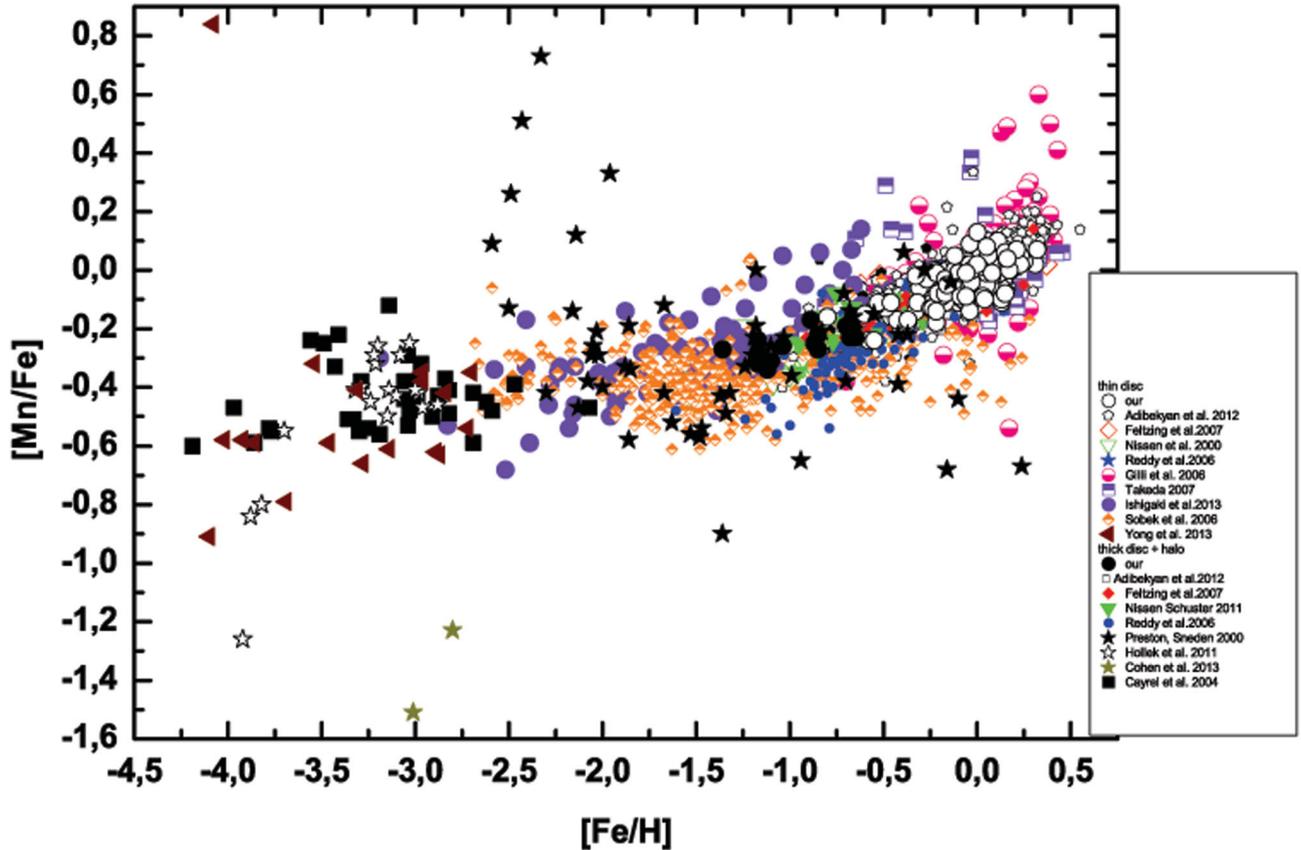


Figure 6. The trend of [Mn/Fe] ratio with respect to [Fe/H] is shown for our stellar sample, in comparison with the data of different authors. Markers are specified in the figure.

this section), but caused by the capture of degenerate electrons with high Fermi energies on protons. Even the reduction of theoretical electron-capture rates (Brachwitz et al. 2000) did not change the amount of ^{55}Co produced in these inner zones, because a region of similar mass content with the relevant Y_e results also in that case. In particular, Yamaguchi et al. (2015) recently have reported the first direct observation of high Mn/Fe ratios in the SNIa remnant 3C 397, that can be explained only by the low Y_e due to electron captures. This makes 3C 397 an ideal candidate of a Chandrasekhar mass SNIa. In addition to these inner zones with normal freeze-out from complete Si-burning, also incomplete Si-burning is taking place in layers further out, where Y_e is determined by the initial metallicity, i.e. where the production of Mn and Fe depends on the initial composition of the stellar progenitor. This results in a situation where the production of Mn and Fe in the inner zones is independent on the initial metallicity (just due to electron capture as a function of central density), while the production of Mn and Fe in the outer zones depends on the initial metallicity (e.g. Nomoto, Thielemann & Yokoi 1984; Thielemann, Nomoto & Yokoi 1986; Iwamoto et al. 1999; Brachwitz et al. 2000; Thielemann et al. 2003; Seitzzahl et al. 2009, 2013a; Fink et al. 2014). If we take, for example, the results of Iwamoto et al. (1999) for the (near-)Chandrasekhar-mass Type Ia explosion model W7 based on progenitor stars of zero and solar metallicity ($[\text{Fe}/\text{H}] = -\infty, 0$), this leads to composition ratios in the ejecta of $[\text{Mn}/\text{Fe}] = 0.067$ and 0.227 , which would of course find their way into the ISM and new stars only after the appropriate delay times for their formation with that initial metallicity. By scaling the Mn production in outer layers varying linearly with metallicity (but keeping constant the Mn yields from the inner ejecta

dominated by electron capture) this would lead for $[\text{Fe}/\text{H}] = 0.25$ and 0.5 to $[\text{Mn}/\text{Fe}] = 0.30$ and 0.38 , respectively. Other delayed detonation models find for solar metallicities values $[\text{Mn}/\text{Fe}] = 0.42$ (see Seitzzahl et al. 2013b). All -mass models lead to $[\text{Mn}/\text{Fe}] > 0$, a contribution needed to explain the change from about -0.4 at low metallicities to $[\text{Mn}/\text{Fe}] > 0.3$ at supersolar metallicities.

Sub-Chandrasekhar-mass Type Ia models (see below) lack the inner electron-capture dominated Mn ejecta and contain only the outer metallicity dependent, incomplete Si-burning ejecta. Therefore, they will eject material with $[\text{Mn}/\text{Fe}] < 0$. The reproduction of the observed [Mn/Fe] trend with respect to [Fe/H] in the Galactic disc is an important diagnostic for GCE and the (Type Ia) supernova models contributing to it. The increasing [Mn/Fe] trend has been considered as a signature: (1) of the gradual enrichment by SNIa ejecta of the ISM (Kobayashi & Nomoto 2009); (2) of the increasing Mn/Fe ratio in the SNIa yields with the metallicity of the progenitor (Cescutti et al. 2008); (3) of the overlapping contribution of sub-Chandrasekhar SNe Ia made by WD mergers (e.g. Pakmor et al. 2010) or triggered by He-detonation on a single WD (e.g. Fink et al. 2010), and SNe Ia reaching the Chandrasekhar mass by accretion on a WD (e.g. Thielemann et al. 1986, 2003; Iwamoto et al. 1999; Brachwitz et al. 2000; Seitzzahl et al. 2009, 2013a; Fink et al. 2014). This last result has been discussed recently by Seitzzahl et al. (2013b): sub-Chandrasekhar SNe Ia do not reach the conditions to make ^{55}Co in nuclear statistical equilibrium opposite to more massive SNe Ia, yielding low Mn/Fe ratio ejecta. Therefore, the [Mn/Fe] trend observed in the Galactic disc may be used as an indirect diagnostic of the relative contribution from different types of SNe Ia. All the three arguments

discussed above may play a role in defining the galactic trend of the [Mn/Fe] ratio, affecting in a similar way the evolution of the [Mn/Fe] with respect to [Fe/H]. The uncertainties associated with the nucleosynthesis of Mn and Fe in SNe Ia need to be also considered. While the nuclear uncertainties seem to be less relevant for this case (Parikh et al. 2013), other uncertainties associated with the SNIa explosion and to the stellar progenitor structure need to be considered.

Mn and Fe have an opposite dependence on the metallicity of the SNIa progenitor, which makes the analysis more complicated. The contribution to the solar inventory by CCSNe and SNe Ia is quite similar for Ni and Fe, yielding a quite flat [Ni/Fe] for stars with metallicities lower than solar in the disc and in the Galactic halo (e.g. Kobayashi et al. 2011; Mishenina et al. 2013, and references therein). As discussed in Mishenina et al. (2013), Ni is a primary product both in CCSNe and in SNe Ia, made by nuclear statistical equilibrium in both the two stellar sources (see Jerkstrand et al. 2015, for a recent analysis of Ni production compared to Fe in CCSN conditions). Therefore, for basic nucleosynthesis reasons the evolution of the [Mn/Ni] ratio should be an observational diagnostic for the production of Mn in stars much easier to study than the [Mn/Fe] ratio.

We need to remind that the reproduction of the observed [Ni/Fe] trend in the galaxy has been proven to be challenging for GCE simulations, both in the halo and in the Galactic disc (e.g. Goswami & Prantzos 2000; Kobayashi et al. 2011). While assumptions made by GCE models may be an important source of uncertainty, the present issues to reproduce the [Ni/Fe] galactic trend is related to the present limitations in theoretical stellar models and, as a consequence, in the stellar yields used by GCE simulations. On the other hand, a confirmation of theoretical results comes for SNe Ia with close to Chandrasekhar mass, with these objects yielding high Ni/Fe and low Mn/Ni ejecta (Yamaguchi et al. 2015). This means that the [Mn/Ni] ratio can be also used to distinguish different SNIa populations together with the [Mn/Fe] ratio, but without being affected by the metallicity dependence associated with the Fe yields of SNe Ia. Consistent observations for Mn, Fe and Ni on the same stellar samples are important to study the production of Fe-group elements in SNe Ia. GCE studies aiming to deliver robust conclusions about the nucleosynthesis of Mn, should take into account both Fe and Ni as reference elements.

In order to study the impact of this in our analysis, we also compare the Mn abundance with Ni (Fig. 5). The average error for [Mn/Ni] is about 0.15 dex (Mishenina et al. 2013). For our stellar sample, the average observed slope for [Mn/Fe] with respect to [Fe/H] is 0.227 ± 0.012 , while we obtain for [Mn/Ni] with respect to [Ni/H] 0.149 ± 0.015 . The two slopes are different. For the considerations made before, we may expect to observe a steeper slope for [Mn/Fe] compared to [Mn/Ni]. This confirm that the metallicity dependence of Fe yields in SNe Ia may play a role in the [Mn/Fe] and [Ni/Fe] trends. The impact of this with respect to the contribution from different SNIa populations to Mn, Fe and Ni still has to be investigated. At the moment we cannot derive any quantitative conclusion, since our thick disc sample does not include enough stars and because of observational errors. Furthermore, in our stellar sample we observe a larger dispersion of the [Mn/Ni] data compared to [Mn/Fe] in Fig. 5, in particular for thin disc stars.

6 CONCLUSIONS AND FINAL REMARKS

In this work, we presented and discussed the abundance measurements of Mn for 247 disc stars. The analysis is based on LTE

assumptions. The insufficient accuracy of atomic data makes difficult to construct an adequate model for NLTE calculations for Mn. We show that in our case the corrections by Bergemann & Gehren (2008) are not confirmed by observations. We have discussed the uncertainties affecting the determination of the Mn abundance. For [Mn/Fe] we obtain an error of about 0.10 dex. The major source of its uncertainty is the choice of the temperature.

For disc stars in our stellar sample we obtain an increasing [Mn/Fe] trend with [Fe/H] consistent with most of other works. Within observational uncertainties we cannot disentangle the abundance patterns for thin disc and thick disc stars, as obtained by Feltzing et al. (2007) and Battistini & Bensby (2015). On the other hand, our determinations of [Mn/Fe] are consistent with the data of these two works for common stars within the observational errors. We have compared the [Mn/Fe] and [Mn/Ni] trends with [Fe/H] and [Ni/H], respectively. The reason is that Mn and Fe production in SNe Ia both depend on the initial metallicity of the progenitor with opposite trends: Mn yields increase with the metallicity of the SNIa progenitor, while Fe yields decrease. On the other hand, Ni production is independent from the initial stellar metallicity. We show that the [Mn/Ni] and [Mn/Fe] patterns have an average slope of 0.149 ± 0.015 and 0.227 ± 0.012 , respectively. While the slopes are different within 2σ , the [Mn/Ni] observational dispersion for thin disc stars and our small sample of thick disc stars do not allow us to derive quantitative conclusions.

We reviewed the production of Mn and Fe in SNe Ia and CCSNe. In particular, there are three main scenarios that are qualitatively compatible with the observed [Mn/Fe] pattern in the Galactic disc, including the relative contribution from both sub-Chandrasekhar mass SNe Ia and more massive SNe Ia. In order to define the relative frequency of the different SNIa populations explaining the [Mn/Fe] observations, the impact of the other two aspects need to be disentangled and weighted consistently by a detailed GCE study.

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REFERENCES

- Adibekyan V. Z., Sousa S. G., Santos N. C., Delgado Mena E., González Hernández J. I., Israelian G., Mayor M., Khachatryan G., 2012, *A&A*, 545, A32
- Allende Prieto C., García López R. J., Lambert D. L., Gustafsson B., 1999, *ApJ*, 527, 879
- Alves-Brito A., Barbuy B., Allen D. M., 2007, in Vazdekis A., Peletier R., eds, *Proc. IAU Symp. 241, Stellar Populations as Building Blocks of Galaxies*. Cambridge Univ. Press, Cambridge, p. 231
- Andrievsky S. M., Korotin S. A., Martin P., 2007, *A&A*, 464, 709
- Baranne A. et al., 1996, *A&AS*, 119, 373
- Battistini C., Bensby T., 2015, *A&A*, 577, A9 (BB 2015)
- Bergemann M., Gehren T., 2007, *A&A*, 473, 291

- Bergemann M., Gehren T., 2008, *A&A*, 492, 823
 Brachwitz F. et al., 2000, *ApJ*, 536, 934
 Bravo E., Domínguez I., Badenes C., Piersanti L., Straniero O., 2010, *ApJ*, 711, L66
 Castelli F., Kurucz R. L., 2004, preprint ([astro-ph/0405087](http://arxiv.org/abs/astro-ph/0405087))
 Cayrel R. et al., 2004, *A&A*, 416, 1117
 Cescutti G., Matteucci F., Lanfranchi G. A., McWilliam A., 2008, *A&A*, 491, 401
 Cohen J. G., Christlieb N., Thompson I., McWilliam A., Shtetman S., Reimers D., Wisotzki L., Kirby E., 2013, *ApJ*, 778, 56
 Cunha K., Smith V. V., Bergemann M., Suntzeff N. B., Lambert D. L., 2010, *ApJ*, 717, 333
 Feltzing S., Fohlman M., Bensby T., 2007, *A&A*, 467, 665
 Fink M., Röpké F. K., Hillebrandt W., Seitzzahl I. R., Sim S. A., Kromer M., 2010, *A&A*, 514, A53
 Fink M. et al., 2014, *MNRAS*, 438, 1762
 Fröhlich C. et al., 2006, *ApJ*, 637, 415
 Galazutdinov G. A., 1992, *Spec. Astrophys. Obs. Russ. Acad. Sci., Crimea*, preprint No. 92
 Gilli G., Israelian G., Ecuivillon A., Santos N. C., Mayor M., 2006, *A&A*, 449, 723
 Goswami A., Prantzos N., 2000, *A&A*, 359, 191
 Gratton R. G., 1989, *A&A*, 208, 171
 Hawkins K., Jofre P., Masseron T., Gilmore G., 2015, *MNRAS*, 453, 758
 Hillebrandt W., Kromer M., Röpké F. K., Ruiter A. J., 2013, *Frontiers Phys.*, 8, 116
 Hix W. R. et al., 2014, *AIP Adv.*, 4, 041013
 Hollek J. K., Frebel A., Roederer I. U., Sneden C., Shetrone M., Beers T. C., Kang S.-J., Thom C., 2011, *ApJ*, 742, 54
 Ishigaki M. N., Aoki W., Chiba M., 2013, *ApJ*, 771, 67
 Iwamoto K., Brachwitz F., Nomoto K., Kishimoto N., Umeda H., Hix W. R., Thielemann F.-K., 1999, *ApJS*, 125, 439
 Jerkstrand A. et al., 2015, *ApJ*, 807, 110
 Katz D., Soubiran C., Cayrel R., Adda M., Cautain R., 1998, *A&A*, 338, 151
 Kobayashi C., Nomoto K., 2009, *ApJ*, 707, 1466
 Kobayashi C., Karakas A. I., Umeda H., 2011, *MNRAS*, 414, 3231
 Kovtyukh V. V., Soubiran C., Belik S. I., Gorlova N. I., 2003, *A&A*, 411, 559
 Kupka F., Piskunov N., Ryabchikova T. A., Stempels H. C., Weiss W. W., 1999, *A&AS*, 138, 119
 Limongi M., Straniero O., Chieffi A., 2000, *ApJS*, 129, 625
 McWilliam A., Rich R. M., Smecker-Hane T. A., 2003, *ApJ*, 592, L21
 Maeda K., Nakamura T., Nomoto K., Mazzali P. A., Patat F., Hachisu I., 2002, *ApJ*, 565, 405
 Matteucci F., Greggio L., 1986, *A&A*, 154, 279
 Matteucci F., Spitoni E., Recchi S., Valiante R., 2009, *A&A*, 501, 531
 Mishenina T. V., Kovtyukh V. V., 2001, *A&A*, 370, 951
 Mishenina T. V., Soubiran C., Kovtyukh V. V., Korotin S. A., 2004, *A&A*, 418, 551
 Mishenina T. V., Soubiran C., Bienayme O., Korotin S. A., Belik S. I., Usenko I. A., Kovtyukh V. V., 2008, *A&A*, 489, 923
 Mishenina T. V., Pignatari M., Korotin S. A., Soubiran C., Charbonnel C., Thielemann F.-K., Gorbaneva T. I., Basak N. Y., 2013, *A&A*, 552, A128
 Nakamura T., Umeda H., Nomoto K., Thielemann F.-K., Burrows A., 1999, *ApJ*, 517, 193
 Nakamura T., Umeda H., Iwamoto K., Nomoto K., Hashimoto M.-A., Hix W. R., Thielemann F.-K., 2001, *ApJ*, 555, 880
 Nissen P. E., Schuster W. J., 2011, *A&A*, 530, A15
 Nissen P. E., Chen Y. Q., Schuster W. J., Zhao G., 2000, *A&A*, 353, 722
 Nomoto K., Thielemann F.-K., Yokoi K., 1984, *ApJ*, 286, 644
 Nomoto K., Tominaga N., Umeda H., Kobayashi C., Maeda K., 2006, *Nucl. Phys. A*, 777, 424
 Nomoto K., Kobayashi C., Tominaga N., 2013, *ARA&A*, 51, 457
 Pakmor R., Kromer M., Röpké F. K., Sim S. A., Ruiter A. J., Hillebrandt W., 2010, *Nature*, 463, 61
 Pancino E., Mucciarelli A., Sbordone L., Bellazzini M., Pasquini L., Monaco L., Ferraro F. R., 2011, *A&A*, 527, A18
 Parikh A., José J., Seitzzahl I. R., Röpké F. K., 2013, *A&A*, 557, A3
 Perruchot S. et al., 2008, in McLean I. S., Casali M. M., eds, *Proc. SPIE Conf. Ser. Vol. 7014, Ground-based and Airborne Instrumentation for Astronomy II*. SPIE, Bellingham, p. 12
 Preston G. W., Sneden C., 2000, *AJ*, 120, 1014
 Prochaska J. X., Naumov S. O., Carney B. W., McWilliam A., Wolfe A. M., 2000, *AJ*, 120, 2513
 Reddy B. E., Lambert D. L., Allende Prieto C., 2006, *MNRAS*, 367, 1329
 Seitzzahl I. R., Meakin C. A., Lamb D. Q., Truran J. W., 2009, *ApJ*, 700, 642
 Seitzzahl I. R. et al., 2013a, *MNRAS*, 429, 1156
 Seitzzahl I. R., Cescutti G., Röpké F. K., Ruiter A. J., Pakmor R., 2013b, *A&A*, 559, L5
 Sobek J. S., Ivans I. I., Simmerer J. A., Sneden C., Hoefflich P., Fulbright J. P., Kraft R. P., 2006, *AJ*, 131, 2949
 Takeda Y., 2007, *PASJ*, 59, 335
 Thielemann F. K., Arnett W. D., 1985, *ApJ*, 295, 604
 Thielemann F.-K., Nomoto K., Yokoi K., 1986, *A&A*, 158, 17
 Thielemann F.-K., Hashimoto M.-A., Nomoto K., 1990, *ApJ*, 349, 222
 Thielemann F.-K., Nomoto K., Hashimoto M.-A., 1996, *ApJ*, 460, 408
 Thielemann F.-K. et al., 2003, in Hillebrandt W., Leibundgut B., eds, *From Twilight to Highlight: The Physics of Supernovae*. Springer-Verlag, Berlin, p. 331
 Thielemann F.-K., Hirschi R., Liebendörfer M., Diehl R., 2011, in Diehl R., Hartmann D. H., Prantzos N., eds, *Lecture Notes in Physics, Vol. 812, Astronomy with Radioactivities*. Springer-Verlag, Berlin, p. 153
 Timmes F. X., Woosley S. E., Weaver T. A., 1995, *ApJS*, 98, 617
 Timmes F. X., Brown E. F., Truran J. W., 2003, *ApJ*, 590, L83
 Travaglio C., Hillebrandt W., Reinecke M., 2005, *A&A*, 443, 1007
 Tymbal V., 1996, in Adelman S. J., Kupka F., Weiss W. W., eds, *ASP Conf. Ser. Vol. 108, M.A.S.S. – Model Atmospheres and Stellar Spectra: 5th Vienna – Workshop*. Astron. Soc. Pac., San Francisco, p. 198
 Umeda H., Nomoto K., 2005, *ApJ*, 619, 427
 Wallerstein G., 1962, *ApJS*, 6, 407
 Wongwathanarat A., Mueller E., Janka H.-T., 2015, *A&A*, 577, A48
 Woosley S. E., Weaver T. A., 1995, *ApJS*, 101, 181
 Woosley S. E., Heger A., Weaver T. A., 2002, *Rev. Mod. Phys.*, 74, 1015
 Yamaguchi H. et al., 2015, *ApJ*, 801, L31
 Yong D. et al., 2013, *ApJ*, 762, 26

SUPPORTING INFORMATION

Additional Supporting Information may be found in the online version of this article:

Table 1. Atmospheric parameters and Mn abundance $\log A(\text{Mn})$ each used line (λ , A) and $[\text{Mn}/\text{Fe}]$ ratio for our target stars (<http://www.mnras.oxfordjournals.org/lookup/suppl/doi:10.1093/mnras/stv2038/-/DC1>).

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